



## Early Journal Content on JSTOR, Free to Anyone in the World

This article is one of nearly 500,000 scholarly works digitized and made freely available to everyone in the world by JSTOR.

Known as the Early Journal Content, this set of works include research articles, news, letters, and other writings published in more than 200 of the oldest leading academic journals. The works date from the mid-seventeenth to the early twentieth centuries.

We encourage people to read and share the Early Journal Content openly and to tell others that this resource exists. People may post this content online or redistribute in any way for non-commercial purposes.

Read more about Early Journal Content at <http://about.jstor.org/participate-jstor/individuals/early-journal-content>.

JSTOR is a digital library of academic journals, books, and primary source objects. JSTOR helps people discover, use, and build upon a wide range of content through a powerful research and teaching platform, and preserves this content for future generations. JSTOR is part of ITHAKA, a not-for-profit organization that also includes Ithaka S+R and Portico. For more information about JSTOR, please contact [support@jstor.org](mailto:support@jstor.org).

that its measured position could be in error by 0.7 of a unit, thus permitting exact coincidence with the krypton line.

It does not follow that the apparent coincidence indicates with any degree of certainty that the principal aurora line is due to krypton. We need only recall the early observations of the principal nebular line, and its apparent coincidence with a prominent nitrogen line, and later with a prominent magnesium line, to be on our guard against conclusions based upon apparent coincidences of single lines in two spectra. In the present case, only one of the several krypton lines seems to have a counterpart in the aurora spectrum.

Again, we should recall that LIVEING and DEWAR produced a strong line at  $\lambda$  5572 by passing a powerful current (spark) of electricity through liquid oxygen and its vapor. This points to the possible oxygen origin of the aurora line.

W. W. CAMPBELL.

#### CORRECTIONS TO WEISSE'S CATALOGUE OF BESSEL STARS.

W. B.<sub>2</sub>. 22<sup>h</sup> 259.

There appears to be an error of about 10'' in declination in the reduction of this star from the time of observation to the epoch of the catalogue. The declination should be increased 10''.

W. B.<sub>2</sub>. 22<sup>h</sup> 395.

The right ascension of this star appears to be 0<sup>s</sup>.6 too large.

W. B.<sub>2</sub>. 22<sup>h</sup> 688.

The declination of this star is in error by 1' 45''. Both the catalogue place and the zone observation should be decreased.

MT. HAMILTON, Dec. 20, 1898.

C. D. PERRINE.

#### COMET *j* 1898 (CHASE).

This comet was discovered November 21, 1898, at the Yale Observatory by Dr. FREDERICK L. CHASE upon photographic plates exposed upon the night of November 14, 1898, for the *Leonid* meteors. Its cometary nature was verified by Dr. CHASE upon the night of November 21, 1898, by making additional exposures.

The comet has also been found upon negatives taken at Harvard College Observatory, at Goodsell Observatory, and at the Lick Observatory.

When discovered, its brightness was estimated to be about 11th magnitude. Theoretically, its maximum brightness was reached about January 1, 1899, and it is now slowly decreasing. By April 1st it will have decreased to about half that at discovery.

From my observations of November 23, December 7, and December 16, 1898, I have computed the following elements for this comet. The first two observations were secured with the 12-inch telescope and the third with the 36-inch refractor:—

*Elements.*

$$\begin{array}{l} T = 1898 \text{ Sept. } 20.15344 \\ \omega = 4^{\circ} \ 37' \ 59''.9 \\ \Omega = 95 \ 51 \ 35.9 \\ i = 22 \ 30 \ 20.3 \end{array} \left. \vphantom{\begin{array}{l} T \\ \omega \\ \Omega \\ i \end{array}} \right\} 1899.0$$

$$\log q = 0.358892$$

A comparison of the ephemeris computed from these elements with observations which I have recently secured with the 36-inch refractor gives the following residuals:—

O—C.

$$\begin{array}{l} \text{Jan. 4, } \Delta \alpha = + 0^s.19, \Delta \delta = + 2''.2 \\ \quad 5, \Delta \alpha = + 0.17, \Delta \delta = + 1.5 \end{array}$$

Soon after the first elements of this comet were published, Professor WEISS, of Vienna, pointed out their similarity to those of Comet 1867 I, and suggested the possibility of an identity. But the residuals given above for dates twenty and twenty-one days after the date of the last observation which was used in computing these elements, show that they are nearly correct, and that the true elements are essentially parabolic. There is, therefore, no probability of the present comet being a return of Comet 1867 I.

MT. HAMILTON, January 16, 1899.

E. F. CODDINGTON.

THE SECOND WASHINGTON STAR CATALOGUE.

This catalogue, which embodies 72,914 observations of stars made with the 8.5-inch transit-circle at the old Naval Observatory between the dates January 3, 1866, and June 30, 1891, was published in November, 1898, as "Appendix I," of the Washington Observations for 1892. It contains the places for 1875.0 of 5151 stars, 185 of which are stars of the American Ephemeris, and the others miscellaneous, ranging in Declination from the North